
pfsspy Documentation

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pfsspy is a python package for carrying out Potential Field Source Surface modelling. For more information on the actually PFSS calculation see [this document](#).

Note: pfsspy is a very new package, so elements of the API are liable to change with the first few releases. If you find any bugs or have any suggestions for improvement, please raise an issue here: <https://github.com/dstansby/pfsspy/issues>

pfsspy can be installed from PyPi using

```
pip install pfsspy
```

Improving performance

1.1 numba

pfsspy automatically detects an installation of `numba`, which compiles some of the numerical code to speed up pfss calculations. To enable this simply `install numba` and use pfsspy as normal.

CHAPTER 2

Citing

If you use pfsspy in work that results in publication, please cite the archived code at *both*

- <https://zenodo.org/record/2566462>
- <https://zenodo.org/record/1472183>

Citation details can be found at the lower right hand of each web page.

For the main user-facing code and a changelog see

3.1 pfsspy Package

3.1.1 Functions

<code>load_output(file)</code>	Load a saved output file.
<code>pfss(input)</code>	Compute PFSS model.

load_output

`pfsspy.load_output(file)`

Load a saved output file.

Loads a file saved using `Output.save()`.

Parameters

file [str, file, Path] File to load.

Returns

Output

pfss

`pfsspy.pfss(input)`

Compute PFSS model.

Extrapolates a 3D PFSS using an eigenfunction method in r, s, p coordinates, on the dumfric grid (equally spaced in $\rho = \ln(r/r_{sun})$, $s = \cos(\theta)$, and $p = \phi$).

The output should have zero current to machine precision, when computed with the DuMFriC staggered discretization.

Parameters

input [*Input*] Input parameters.

Returns

out [*Output*]

3.1.2 Classes

<i>FieldLine</i> (x, y, z, dtime, output)	A single magnetic field line.
<i>Grid</i> (ns, nphi, nr, rss)	Grid on which the solution is calculated.
<i>Input</i> (br, nr, rss[, dtime])	Input to PFSS modelling.
<i>Output</i> (alr, als, alp, grid[, dtime])	Output of PFSS modelling.

FieldLine

class pfsspy.**FieldLine** (x, y, z, dtime, output)

Bases: `object`

A single magnetic field line.

Parameters

x, y, z [array] Field line coordinates in a Carrington frame of reference. Must be in units of solar radii.

dtime [astropy.time.Time] Time at which the field line was traced. Needed for transforming the field line coordinates to other coordinate frames.

output [Output] The PFSS output through which this field line was traced.

Attributes

coords [astropy.coordinates.SkyCoord] Field line coordinates.

Attributes Summary

<i>expansion_factor</i>	Magnetic field expansion factor.
<i>footpoints</i>	Magnetic field footpoint(s).
<i>is_open</i>	Returns <code>True</code> if one of the field line is connected to the solar surface and one to the outer boundary, <code>False</code> otherwise.
<i>polarity</i>	Magnetic field line polarity.

Attributes Documentation

expansion_factor

Magnetic field expansion factor.

The expansion factor is defined as $(r_{\odot}^2 B_{\odot}) / (r_{ss}^2 B_{ss})$

Returns

exp_fact [float] Field line expansion factor. If field line is closed, returns np.nan.

footpoints

Magnetic field footpoint(s).

These are the ends of the magnetic field line that lie on the solar surface. For an open field line there is one footpoint, for a close field line there are two footpoints.

Returns

footpoints [[SkyCoord](#)]

is_open

Returns `True` if one of the field line is connected to the solar surface and one to the outer boundary, `False` otherwise.

polarity

Magnetic field line polarity.

Returns

pol [int] 0 if the field line is closed, otherwise $\text{sign}(B_r)$ of the magnetic field on the solar surface.

Grid

class `pfsspy.Grid`(*ns, nphi, nr, rss*)

Bases: `object`

Grid on which the solution is calculated.

The grid is evenly spaced in ($\cos(\theta)$, ϕ , $\log(r)$). See `pfsspy.coords` for more information.

Attributes Summary

<i>dp</i>	Cell size in ϕ .
<i>dr</i>	Cell size in $\log(r)$.
<i>ds</i>	Cell size in $\cos(\theta)$.
<i>pc</i>	Location of the centre of cells in ϕ .
<i>pg</i>	Location of the edges of grid cells in ϕ .
<i>rc</i>	Location of the centre of cells in $\log(r)$.
<i>rg</i>	Location of the edges of grid cells in $\log(r)$.
<i>sc</i>	Location of the centre of cells in $\cos(\theta)$.
<i>sg</i>	Location of the edges of grid cells in $\cos(\theta)$.

Attributes Documentation

dp

Cell size in ϕ .

dr

Cell size in $\log(r)$.

ds

Cell size in $\cos(\theta)$.

pc

Location of the centre of cells in ϕ .

pg Location of the edges of grid cells in phi.

rc Location of the centre of cells in log(r).

rg Location of the edges of grid cells in log(r).

sc Location of the centre of cells in cos(theta).

sg Location of the edges of grid cells in cos(theta).

Input

class pfsspy.Input (*br, nr, rss, dtime=None*)

Bases: `object`

Input to PFSS modelling.

Warning: The input must be on a regularly spaced grid in ϕ and $s = \cos(\theta)$. See [pfsspy.coords](#) for more information on the coordinate system.

Parameters

br [2D array, `sunpy.map.Map`] Boundary condition of radial magnetic field at the inner surface. If a SunPy map is automatically extracted as `map.data` with *no* processing.

nr [int] Number of cells in the radial direction to calculate the PFSS solution on.

rss [float] Radius of the source surface, as a fraction of the solar radius.

dtime [datetime, optional] Datetime at which the input map was measured. If given it is attached to the output and any field lines traced from the output.

Methods Summary

<code>plot_input(self, ax)</code>	Plot a 2D image of the magnetic field boundary condition.
-----------------------------------	---

Methods Documentation

plot_input (*self, ax=None, **kwargs*)

Plot a 2D image of the magnetic field boundary condition.

Parameters

ax [Axes] Axes to plot to. If `None`, creates a new figure.

Output

class pfsspy.Output (*alr, als, alp, grid, dtime=None*)

Bases: `object`

Output of PFSS modelling.

Parameters

- alr** : Vector potential * grid spacing in radial direction.
- als** : Vector potential * grid spacing in elevation direction.
- alp** : Vector potential * grid spacing in azimuth direction.
- grid** [Grid] Grid that the output was calculated on.
- dtime** [datetime, optional] Datetime at which the input was measured.

Attributes Summary

<i>al</i>	Vector potential times cell edge lengths.
<i>bc</i>	B on the centres of the cell faces.
<i>bg</i>	B as a (weighted) averaged on grid points.
<i>source_surface_br</i>	Br on the source surface.

Methods Summary

<i>plot_pil</i> (self[, ax])	Plot the polarity inversion line on the source surface.
<i>plot_source_surface</i> (self[, ax])	Plot a 2D image of the magnetic field at the source surface.
<i>save</i> (self, file)	Save the output to file.
<i>trace</i> (self, tracer, seeds)	

Parameters

Attributes Documentation

al

Vector potential times cell edge lengths.
Returns $ar * L_r$, $as * L_s$, $ap * L_p$ on cell edges.

bc

B on the centres of the cell faces.

bg

B as a (weighted) averaged on grid points.

Returns

br [array]

bs [array]

bp [array]

source_surface_br

Br on the source surface.

Methods Documentation

plot_pil (*self*, *ax=None*, ***kwargs*)

Plot the polarity inversion line on the source surface.

The PIL is where $B_r = 0$.

Parameters

ax [Axes] Axes to plot to. If *None*, creates a new figure.

****kwargs** : Keyword arguments are handed to *ax.contour*.

plot_source_surface (*self*, *ax=None*, ***kwargs*)

Plot a 2D image of the magnetic field at the source surface.

Parameters

ax [Axes] Axes to plot to. If *None*, creates a new figure.

kwargs : Additional keyword arguments are handed to *pcolormesh* that renders the source surface. A useful option here is handing *rasterized=True* to rasterize the image.

save (*self*, *file*)

Save the output to file.

This saves the required information to reconstruct an Output object in a compressed binary numpy file (see `numpy.savez_compressed()` for more information). The file extension is *.npz*, and is automatically added if not present.

Parameters

file [str, file, [Path](#)] File to save to. If *.npz* extension isn't present it is added when saving the file.

trace (*self*, *tracer*, *seeds*)

Parameters

tracer [[tracing.Tracer](#)]

seeds [(n, 3) shaped array] Starting coordinates, in cartesian coordinates. [pfsspy.coords](#) can be used to convert from spherical coordinates to cartesian coordinates and vice versa.

3.2 pfsspy.tracing Module

3.2.1 Classes

PythonTracer ([<i>atol</i> , <i>rtol</i>])	Tracer using native python code.
Tracer	Abstract base class for a streamline tracer.

PythonTracer

class `pfsspy.tracing.PythonTracer` (*atol=0.0001*, *rtol=0.0001*)

Bases: [pfsspy.tracing.Tracer](#)

Tracer using native python code.

Uses `scipy.integrate.solve_ivp`, with an LSODA method.

Methods Summary

`trace(self, seeds, output)`

Parameters

Methods Documentation

`trace (self, seeds, output)`

Parameters

seeds [(n, 3) array] Coordinates of the magnetic field seed points.

output [pfsspy.Output] pfss output.

Returns

streamlines [list of FieldLine] List of traced streamlines

Tracer

class pfsspy.tracing.Tracer

Bases: `abc.ABC`

Abstract base class for a streamline tracer.

Methods Summary

`cartesian_to_coordinate()`

Convert cartesian coordinate outputted by a tracer to a *FieldLine* object.

`trace(self, seeds, output)`

Parameters

`validate_seeds_shape(seeds)`

Check that *seeds* has the right shape.

Methods Documentation

static `cartesian_to_coordinate()`

Convert cartesian coordinate outputted by a tracer to a *FieldLine* object.

trace (self, seeds, output)

Parameters

seeds [(n, 3) array] Coordinates of the magnetic field seed points.

output [pfsspy.Output] pfss output.

Returns

streamlines [list of FieldLine] List of traced streamlines

static `validate_seeds_shape (seeds)`

Check that *seeds* has the right shape.

3.3 Changelog

3.3.1 0.3.0

- The API for doing magnetic field tracing has changed. The new `pfsspy.tracing` module contains `Tracer` classes that are used to perform the tracing. Code needs to be changed from:

```
fline = output.trace(x0)
```

to:

```
tracer = pfsspy.tracing.PythonTracer()
tracer.trace(x0, output)
flines = tracer.xs
```

Additionally `x0` can be a 2D array that contains multiple seed points to trace, taking advantage of the parallelism of some solvers.

- The `pfsspy.FieldLine` class no longer inherits from `SkyCoord`, but the `SkyCoord` coordinates are now stored in `pfsspy.FieldLine.coords` attribute.
- `pfsspy.FieldLine.expansion_factor` now returns `np.nan` instead of `None` if the field line is closed.
- `pfsspy.FieldLine` now has a `footpoints` attribute that returns the footpoint(s) of the field line.

3.3.2 0.2.0

- `pfsspy.Input` and `pfsspy.Output` now take the optional keyword argument `dtime`, which stores the datetime on which the magnetic field measurements were made. This is then propagated to the `obstime` attribute of computed field lines, allowing them to be transformed in to coordinate systems other than Carrington frames.
- `pfsspy.FieldLine` no longer overrides the `SkyCoord __init__`; this should not matter to users, as `FieldLine` objects are constructed internally by calling `pfsspy.Output.trace()`

3.3.3 0.1.5

- `Output.plot_source_surface` now accepts keyword arguments that are given to Matplotlib to control the plotting of the source surface.

3.3.4 0.1.4

- Added more explanatory comments to the examples
- Corrected the dipole solution calculation
- Added `pfsspy.coords.sph2cart()` to transform from spherical to cartesian coordinates.

3.3.5 0.1.3

- `pfsspy.Output.plot_pil()` now accepts keyword arguments that are given to Matplotlib to control the style of the contour.

- `pfsspy.FieldLine.expansion_factor` is now cached, and is only calculated once if accessed multiple times.

for usage examples see

3.4 pfsspy examples

Note: Click [here](#) to download the full example code

3.4.1 Dipole source solution

A simple example showing how to use pfsspy to compute the solution to a dipole source field.

First, import required modules

```
import astropy.constants as const
import matplotlib.pyplot as plt
import matplotlib.patches as mpatch
import numpy as np
import pfsspy
import pfsspy.coords as coords
```

To start with we need to construct an input for the PFSS model. To do this, first set up a regular 2D grid in (phi, s), where $s = \cos(\theta)$ and (phi, theta) are the standard spherical coordinate system angular coordinates. In this case the resolution is (360 x 180).

```
nphi = 360
ns = 180

phi = np.linspace(0, 2 * np.pi, nphi)
s = np.linspace(-1, 1, ns)
s, phi = np.meshgrid(s, phi)
```

Now we can take the grid and calculate the boundary condition magnetic field.

```
def dipole_Br(r, s):
    return 2 * s / r**3

br = dipole_Br(1, s).T
```

The PFSS solution is calculated on a regular 3D grid in (phi, s, rho), where $\rho = \ln(r)$, and r is the standard spherical radial coordinate. We need to define the number of rho grid points, and the source surface radius.

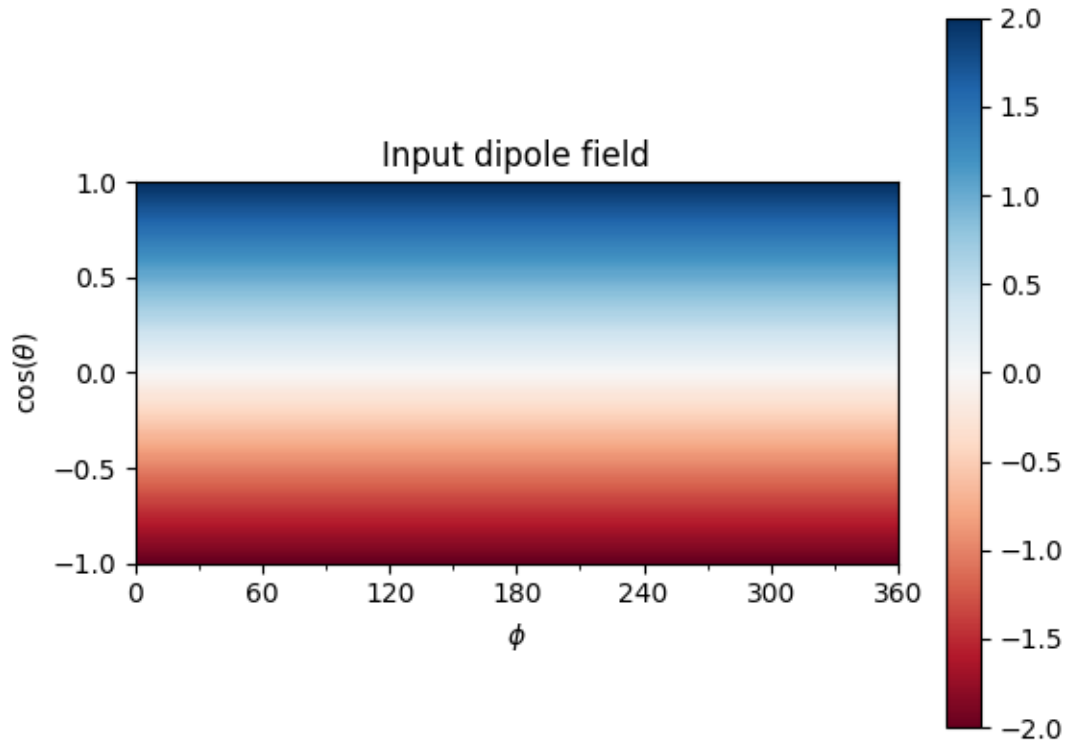
```
nrho = 50
rss = 2.5
```

From the boundary condition, number of radial grid points, and source surface, we now construct an Input object that stores this information

```
input = pfsspy.Input(br, nrho, rss)
```

Using the Input object, plot the input field

```
fig, ax = plt.subplots()
mesh = input.plot_input(ax)
fig.colorbar(mesh)
ax.set_title('Input dipole field')
```

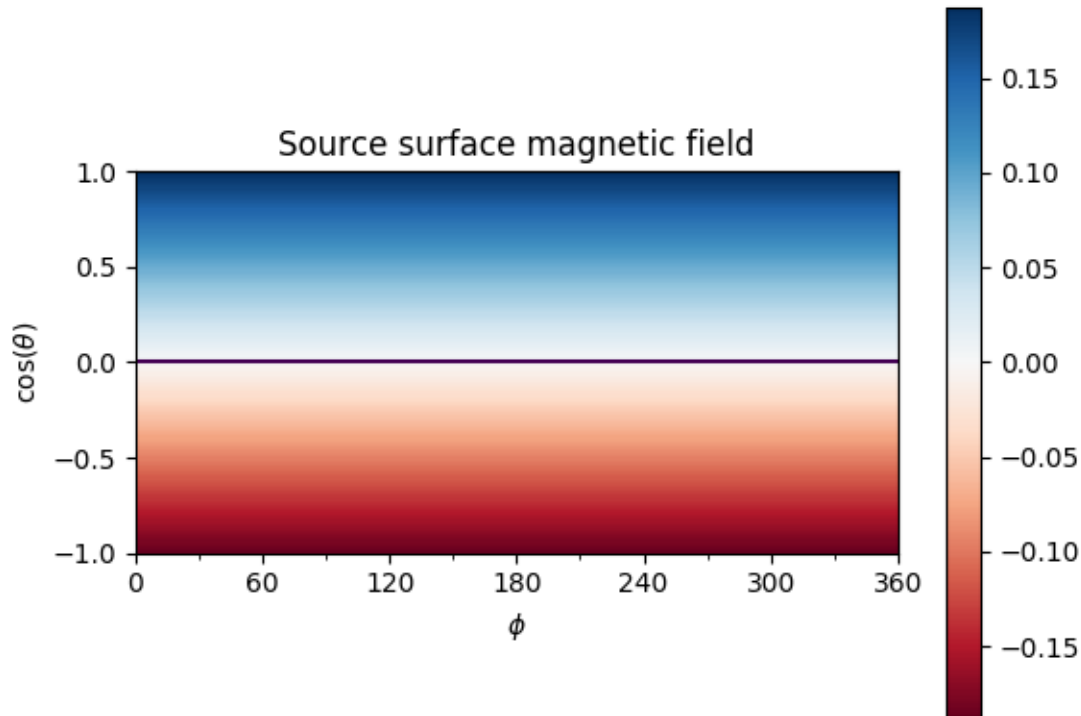


Now calculate the PFSS solution.

```
output = pfsspy.pfss(input)
```

Using the Output object we can plot the source surface field, and the polarity inversion line.

```
fig, ax = plt.subplots()
mesh = output.plot_source_surface(ax)
fig.colorbar(mesh)
output.plot_pil(ax)
ax.set_title('Source surface magnetic field')
```



Finally, using the 3D magnetic field solution we can trace some field lines. In this case 32 points equally spaced in theta are chosen and traced from the source surface outwards.

```
fig, ax = plt.subplots()
ax.set_aspect('equal')

# Take 32 start points spaced equally in theta
r = 1.01
phi = np.pi / 2
r = 1.01
phi = np.pi / 2
theta = np.linspace(0, np.pi, 33)
x0 = np.array(coords.sph2cart(r, theta, phi)).T

tracer = pfsspy.tracing.PythonTracer()
field_lines = tracer.trace(x0, output)

for field_line in field_lines:
    color = {0: 'black', -1: 'tab:blue', 1: 'tab:red'}.get(field_line.polarity)
    ax.plot(field_line.coords.y / const.R_sun,
            field_line.coords.z / const.R_sun, color=color)

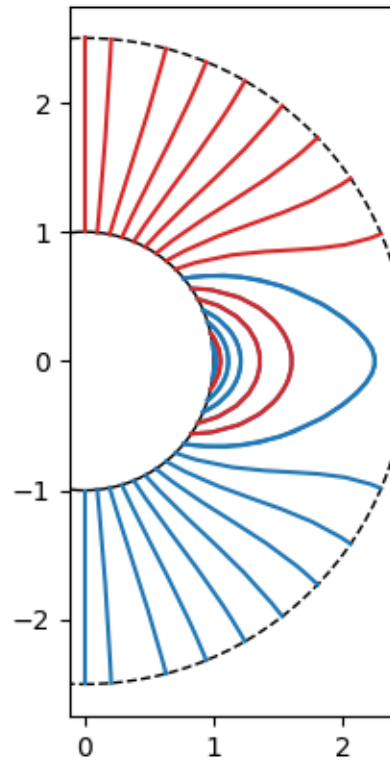
# Add inner and outer boundary circles
ax.add_patch(mpatch.Circle((0, 0), 1, color='k', fill=False))
ax.add_patch(mpatch.Circle((0, 0), input.grid.rss, color='k', linestyle='--',
                           fill=False))
```

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```
ax.set_title('PFSS solution for a dipole source field')
plt.show()
```

PFSS solution for a dipole source field



Total running time of the script: (0 minutes 27.265 seconds)

Note: Click [here](#) to download the full example code

3.4.2 GONG PFSS extrapolation

Calculating PFSS solution for a GONG synoptic magnetic field map.

First, import required modules

```
import os
import astropy.constants as const
import matplotlib.pyplot as plt
from mpl_toolkits.mplot3d import Axes3D
import numpy as np
import sunpy.map

import pfsspy
```

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```
from pfsspy import coords
from pfsspy import tracing
```

Load a GONG magnetic field map. If 'gong.fits' is present in the current directory, just use that, otherwise download a sample GONG map.

```
if not os.path.exists('190310t0014gong.fits') and not os.path.exists('190310t0014gong.
↳fits.gz'):
    import urllib.request
    urllib.request.urlretrieve(
        'https://gong2.nso.edu/oQR/zqs/201903/mrzqs190310/mrzqs190310t0014c2215_333.
↳fits.gz',
        '190310t0014gong.fits.gz')

if not os.path.exists('190310t0014gong.fits'):
    import gzip
    with gzip.open('190310t0014gong.fits.gz', 'rb') as f:
        with open('190310t0014gong.fits', 'wb') as g:
            g.write(f.read())
```

We can now use SunPy to load the GONG fits file, and extract the magnetic field data.

The mean is subtracted to enforce $\text{div}(\mathbf{B}) = 0$ on the solar surface: n.b. it is not obvious this is the correct way to do this, so use the following lines at your own risk!

```
[[br, header]] = sunpy.io.fits.read('190310t0014gong.fits')
br = br - np.mean(br)
```

GONG maps have their LH edge at -180deg in Carrington Longitude, so roll to get it at 0deg. This way the input magnetic field is in a Carrington frame of reference, which matters later when lining the field lines up with the AIA image.

```
br = np.roll(br, header['CRVAL1'] + 180, axis=1)
```

The PFSS solution is calculated on a regular 3D grid in (phi, s, rho), where $\rho = \ln(r)$, and r is the standard spherical radial coordinate. We need to define the number of rho grid points, and the source surface radius.

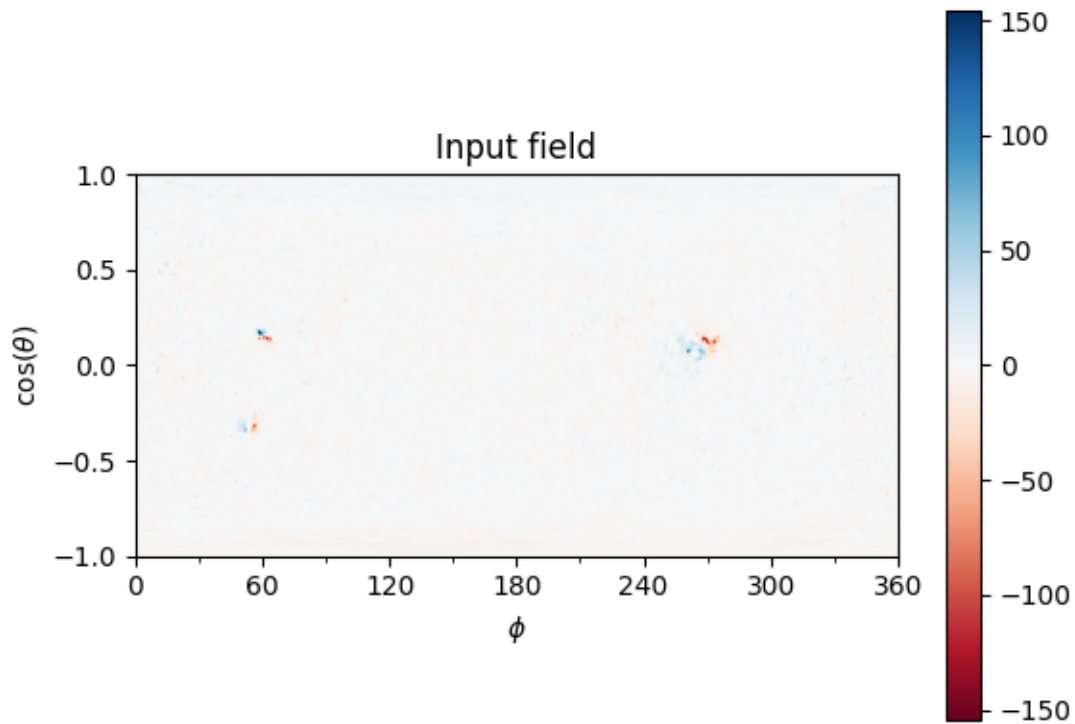
```
nrho = 60
rss = 2.5
```

From the boundary condition, number of radial grid points, and source surface, we now construct an Input object that stores this information

```
input = pfsspy.Input(br, nrho, rss)
```

Using the Input object, plot the input field

```
fig, ax = plt.subplots()
mesh = input.plot_input(ax)
fig.colorbar(mesh)
ax.set_title('Input field')
```

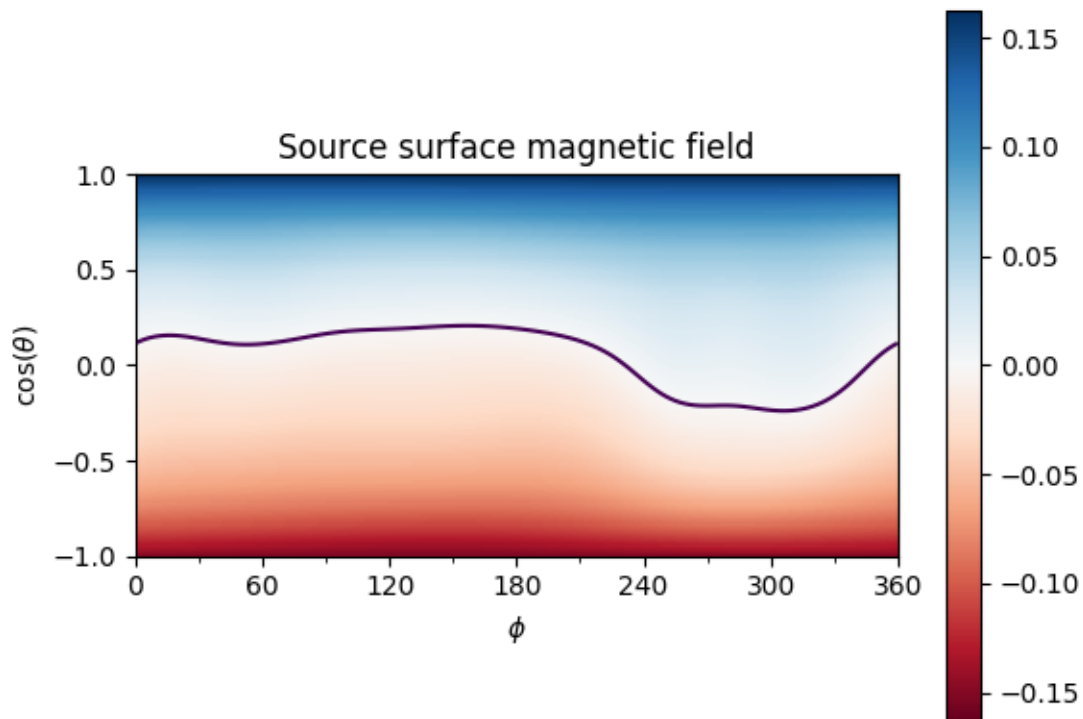


Now calculate the PFSS solution, and plot the polarity inversion line.

```
output = pfsspy.pfss(input)
output.plot_pil(ax)
```

Using the Output object we can plot the source surface field, and the polarity inversion line.

```
fig, ax = plt.subplots()
mesh = output.plot_source_surface(ax)
fig.colorbar(mesh)
output.plot_pil(ax)
ax.set_title('Source surface magnetic field')
```



Finally, using the 3D magnetic field solution we can trace some field lines. In this case 256 points equally gridded in theta and phi are chosen and traced from the source surface outwards.

```
fig = plt.figure()
ax = fig.add_subplot(111, projection='3d')

tracer = tracing.PythonTracer()
# Loop through 16 values in theta and 16 values in phi
r = 1.01
theta = np.linspace(0, np.pi, 17)
phi = np.linspace(0, 2 * np.pi, 17)
theta, phi = np.meshgrid(theta, phi)
theta, phi = theta.ravel(), phi.ravel()

seeds = np.array(coords.sph2cart(r, theta, phi)).T

field_lines = tracer.trace(seeds, output)

for field_line in field_lines:
    color = {0: 'black', -1: 'tab:blue', 1: 'tab:red'}.get(field_line.polarity)
    ax.plot(field_line.coords.x / const.R_sun,
            field_line.coords.y / const.R_sun,
            field_line.coords.z / const.R_sun,
            color=color, linewidth=1)

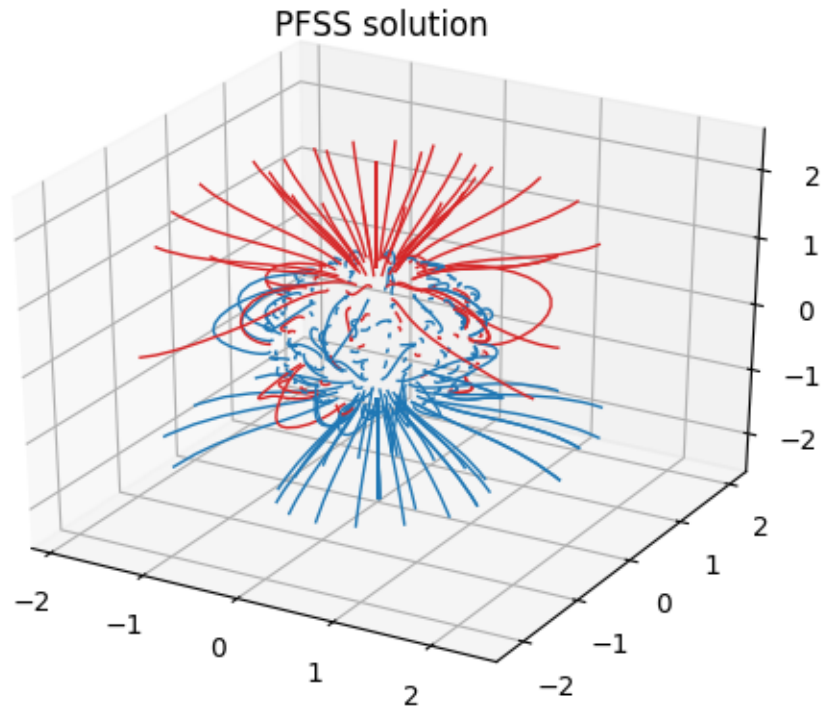
ax.set_title('PFSS solution')
```

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```
plt.show()

# sphinx_gallery_thumbnail_number = 3
```



Total running time of the script: (1 minutes 13.381 seconds)

Note: Click [here](#) to download the full example code

3.4.3 Overplotting field lines on AIA maps

This example shows how to take a PFSS solution, trace some field lines, and overplot the traced field lines on an AIA 193 map.

First, we import the required modules

```
from datetime import datetime
import os

import astropy.units as u
import matplotlib.pyplot as plt
import numpy as np
import sunpy.map
```

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```
import sunpy.io.fits

import pfsspy
import pfsspy.coords as coords
import pfsspy.tracing as tracing
```

Load a GONG magnetic field map. The map date is 10/03/2019

```
if not os.path.exists('190310t0014gong.fits') and not os.path.exists('190310t0014gong.
↳fits.gz'):
    import urllib.request
    urllib.request.urlretrieve(
        'https://gong2.nso.edu/oQR/zqs/201903/mrzqs190310/mrzqs190310t0014c2215_333.
↳fits.gz',
        '190310t0014gong.fits.gz')

if not os.path.exists('190310t0014gong.fits'):
    import gzip
    with gzip.open('190310t0014gong.fits.gz', 'rb') as f:
        with open('190310t0014gong.fits', 'wb') as g:
            g.write(f.read())
```

Load the corresponding AIA 193 map

```
if not os.path.exists('AIA20190310.fits'):
    import urllib.request
    urllib.request.urlretrieve(
        'http://jsoc2.stanford.edu/data/aia/synoptic/2019/03/10/H0000/AIA20190310_
↳0000_0193.fits',
        'AIA20190310.fits')

aia = sunpy.map.Map('AIA20190310.fits')
datetime = aia.date
```

We can now use SunPy to load the GONG fits file, and extract the magnetic field data.

The mean is subtracted to enforce $\text{div}(\mathbf{B}) = 0$ on the solar surface: n.b. it is not obvious this is the correct way to do this, so use the following lines at your own risk!

```
[[br, header]] = sunpy.io.fits.read('190310t0014gong.fits')
br = br - np.mean(br)
```

GONG maps have their LH edge at -180deg in Carrington Longitude, so roll to get it at 0deg. This way the input magnetic field is in a Carrington frame of reference, which matters later when lining the field lines up with the AIA image.

```
br = np.roll(br, header['CRVAL1'] + 180, axis=1)
```

The PFSS solution is calculated on a regular 3D grid in (ϕ, s, ρ) , where $\rho = \ln(r)$, and r is the standard spherical radial coordinate. We need to define the number of grid points in ρ , and the source surface radius.

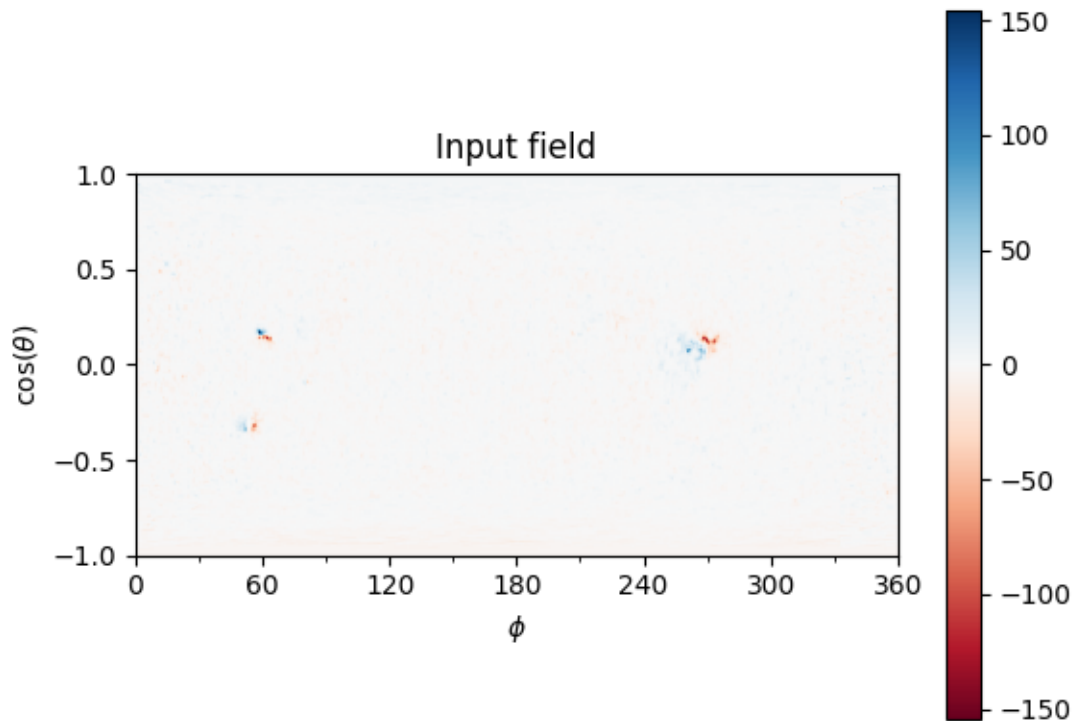
```
nrho = 60
rss = 2.5
```

From the boundary condition, number of radial grid points, and source surface, we now construct an *Input* object that stores this information

```
input = pfsspy.Input(br, nrho, rss, dtype=dtime)
```

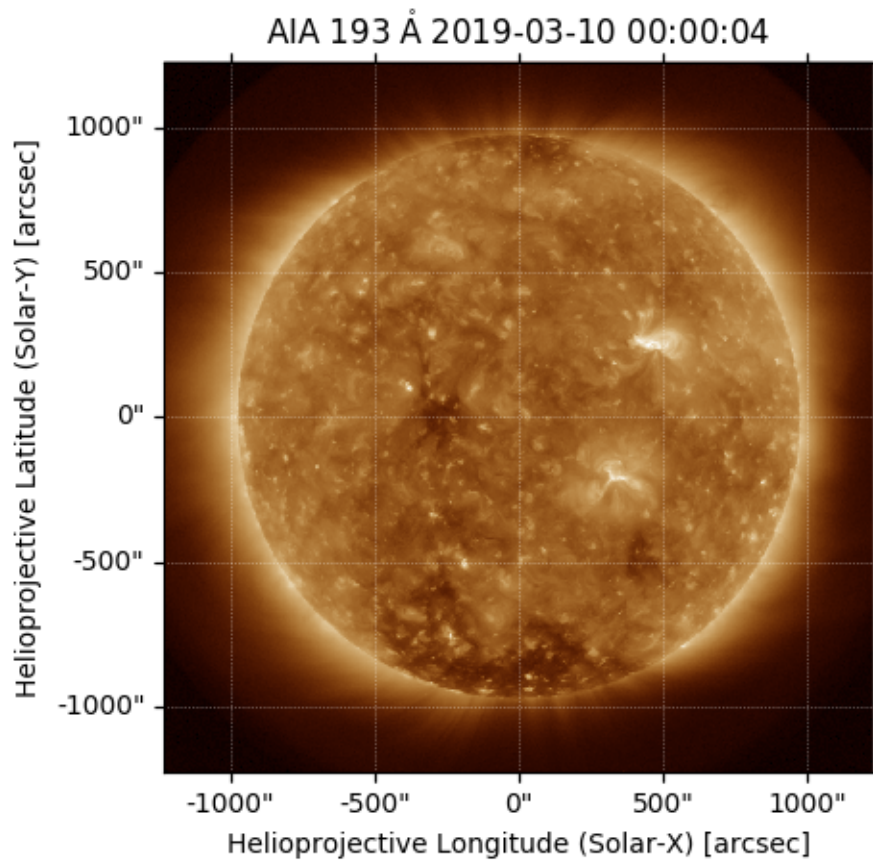
Using the *Input* object, plot the input photospheric magnetic field

```
fig, ax = plt.subplots()
mesh = input.plot_input(ax)
fig.colorbar(mesh)
ax.set_title('Input field')
```



We can also plot the AIA map to give an idea of the global picture. There is a nice active region in the top right of the AIA plot, that can also be seen in the top left of the photospheric field plot above.

```
ax = plt.subplot(1, 1, 1, projection=aia)
aia.plot(ax)
```



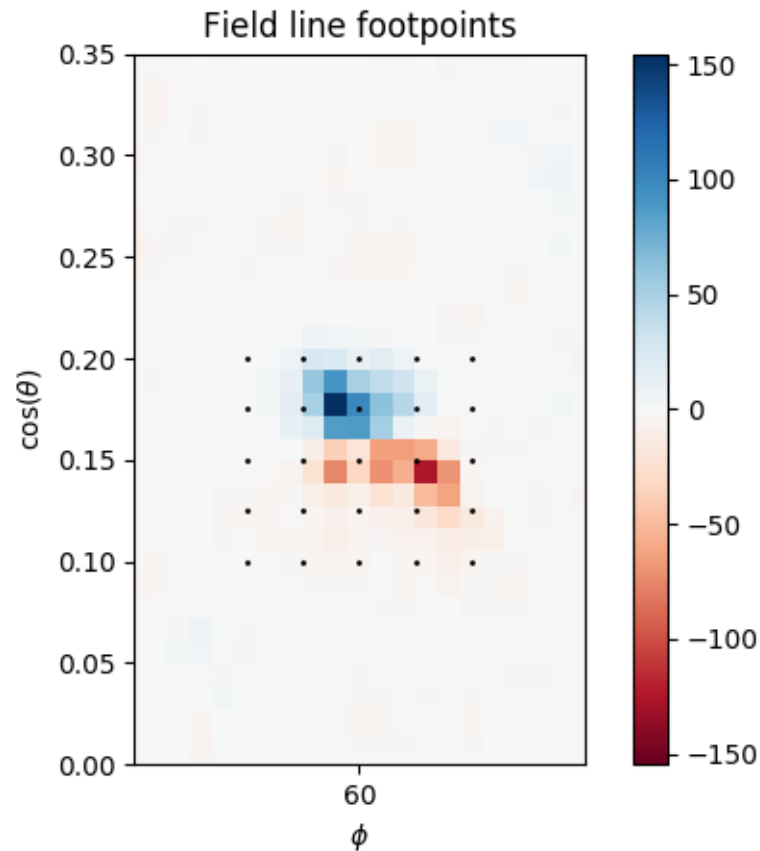
Now we construct a 10 x 10 grid of footpoints to trace some magnetic field lines from.

The figure shows a zoom in of the magnetic field map, with the footpoints overplotted. The footpoints are centered around the active region mentioned above.

```
s, phi = np.meshgrid(np.linspace(0.1, 0.2, 5),
                     np.deg2rad(np.linspace(55, 65, 5)))

fig, ax = plt.subplots()
mesh = input.plot_input(ax)
fig.colorbar(mesh)
ax.scatter(np.rad2deg(phi), s, color='k', s=1)

ax.set_xlim(50, 70)
ax.set_ylim(0, 0.35)
ax.set_title('Field line footpoints')
```



Compute the PFSS solution from the GONG magnetic field input

```
output = pfsspy.pfss(input)
```

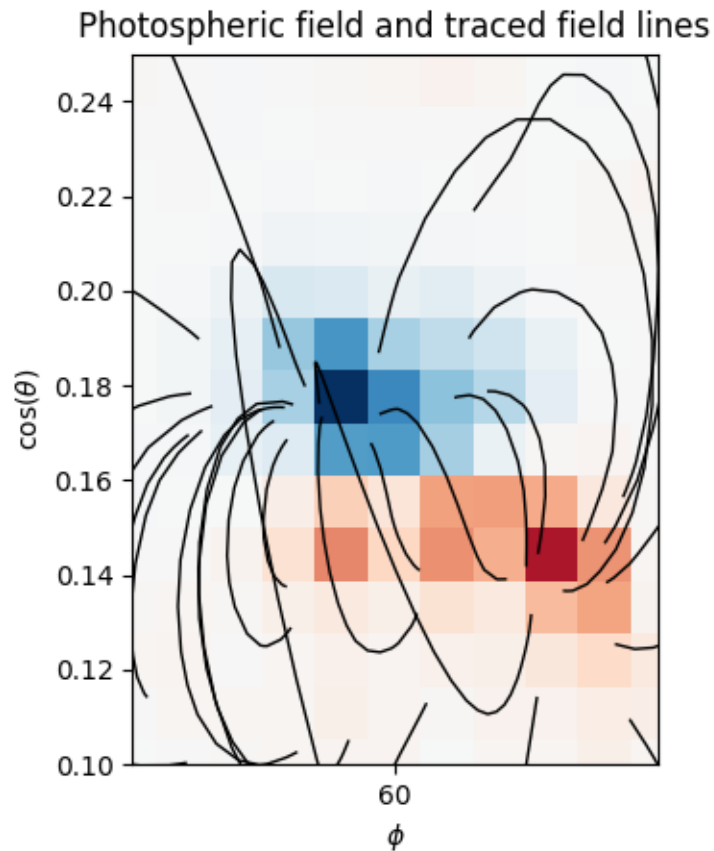
Trace field lines from the footpoints defined above. `pfsspy.coords` is used to convert the s , ϕ coordinates into the cartesian coordinates that are needed by the tracer.

```
tracer = tracing.PythonTracer(atol=1e-8)
flines = []
for s, phi in zip(s.ravel(), phi.ravel()):
    x0 = np.array(pfsspy.coords.strum2cart(0.01, s, phi))
    flines += tracer.trace(x0, output)
```

Plot the input GONG magnetic field map, along with the traced magnetic field lines.

```
fig, ax = plt.subplots()
mesh = input.plot_input(ax)
for fline in flines:
    fline.coords.representation_type = 'spherical'
    ax.plot(fline.coords.lon / u.deg, np.sin(fline.coords.lat),
            color='black', linewidth=1)

ax.set_xlim(55, 65)
ax.set_ylim(0.1, 0.25)
ax.set_title('Photospheric field and traced field lines')
```

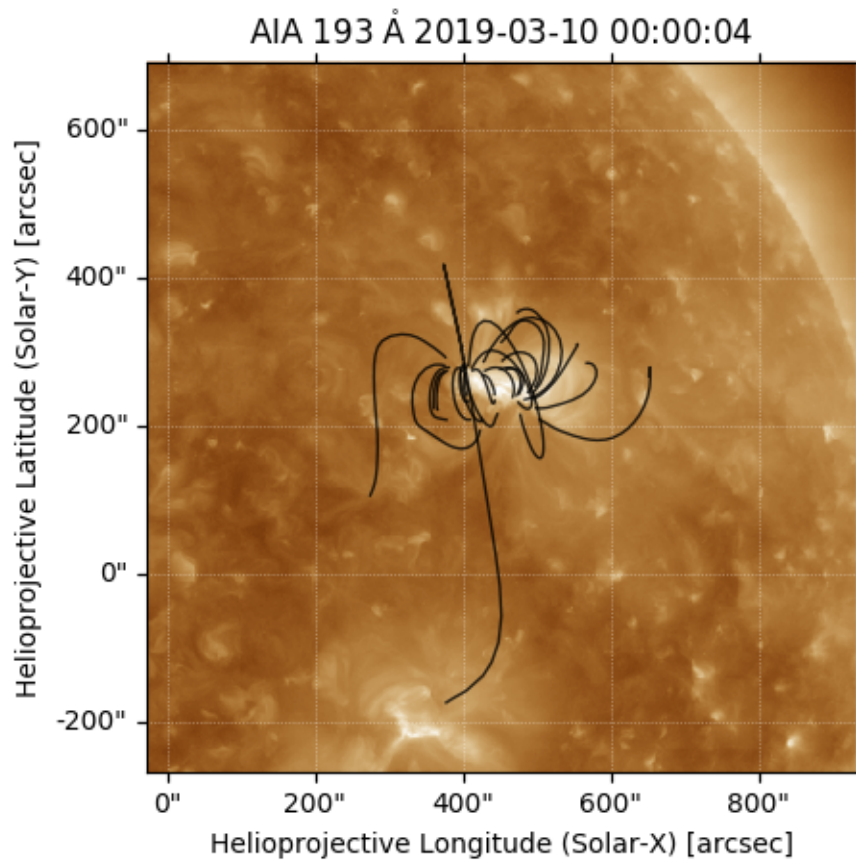


Plot the AIA map, along with the traced magnetic field lines. Inside the loop the field lines are converted to the AIA observer coordinate frame, and then plotted on top of the map.

```
fig = plt.figure()
ax = plt.subplot(1, 1, 1, projection=aia)
transform = ax.get_transform('world')
aia.plot(ax)
for fline in flines:
    coords = fline.coords.transform_to(aia.coordinate_frame)
    ax.plot_coord(coords, alpha=0.8, linewidth=1, color='black')

ax.set_xlim(500, 900)
ax.set_ylim(400, 800)
plt.show()

# sphinx_gallery_thumbnail_number = 5
```



Total running time of the script: (0 minutes 55.699 seconds)

and for the helper modules (behind the scene!) see

3.5 Helper modules

3.5.1 pfsspy.plot Module

Functions

`contour(phi, costheta, field, levels[, ax])`

Parameters

`radial_cut(phi, costheta, field[, ax])`

contour

`pfsspy.plot.contour(phi, costheta, field, levels, ax=None, **kwargs)`

Parameters

phi :

costheta :
field :
levels :
ax [Axes, optional] Axes to plot to. If None a new figure is created.
****kwargs :** Keyword arguments are handed to *ax.contour*.

radial_cut

`pfsspy.plot.radial_cut(phi, costheta, field, ax=None, **kwargs)`

3.5.2 pfsspy.coords Module

Helper functions for coordinate transformations used in the PFSS domain.

The PFSS solution is calculated on a “strumfric” grid defined by

- $\rho = \log(r)$
- $s = \cos(\theta)$
- ϕ

where r, θ, ϕ are spherical coordinates that have ranges

- $1 < r < r_{ss}$
- $0 < \theta < \pi$
- $0 < \phi < 2\pi$

The transformation between cartesian coordinates used by the tracer and the above coordinates is given by

- $x = r \sin(\theta) \cos(\phi)$
- $y = r \sin(\theta) \sin(\phi)$
- $z = r \cos(\theta)$

Functions

<code>cart2strum(x, y, z)</code>	Convert cartesian coordinates to strumfric coordinates.
<code>sph2cart(r, theta, phi)</code>	Convert spherical coordinates to cartesian coordinates.
<code>strum2cart(rho, s, phi)</code>	Convert strumfric coordinates to cartesian coordinates.

cart2strum

`pfsspy.coords.cart2strum(x, y, z)`
 Convert cartesian coordinates to strumfric coordinates.

sph2cart

`pfsspy.coords.sph2cart(r, theta, phi)`
 Convert spherical coordinates to cartesian coordinates.

strum2cart

`pfsspy.coords.strum2cart` (*rho, s, phi*)

Convert strumfric coordinates to cartesian coordinates.

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